

# Geochemical constraints on volatile delivery to the inner solar system

*Comets and LHB, Gdynia June 2014*

Uffe Gråe Jørgensen, Copenhagen University, Denmark

Niels Bohr Institute & Centre for Star and Planet Formation

**Did water on Earth and Mars condense from the protosolar nebula together with the rest of the material that made the planets,**

**or did it come with the late veneer objects**

**or was Late Heavy Bombardment (LHB) the source ????**

No single source seems to be able to explain all the main observables,

so a combination of objects is most likely necessary;

The question is therefore "which sources and in which proportion?"

How can a combined study of Earth, Moon, and Mars help?

**Isua tells us that the Earth's crust was not marked by a chondritic bombardment 3.8 Gyr ago (the time of the Isua sedimentation).**

# **Earth's water is not from the original nebula condensation.....**

Core accretion models and observations of protostellar disks do not support condensation of water at 1 AU around solar type stars.

Asteroids inside 2.7 AU are dehydrated (e.g. S-type, E-type, ordinary chondrites), outside they may be hydrated (e.g. carbonaceous chondrites with 1% - 10% water, representing C-type asteroids).

In the inner solar system, material forming the planets accreted locally (as seen e.g. by the distinct oxygen isotopes)

**Therefore the Earth most likely formed dry,  
and water must have been delivered later,  
during the late veneer (4.5-3.9 Gyr ago?)  
or during the LHB (3.9-3.8 Gyr ago?)**

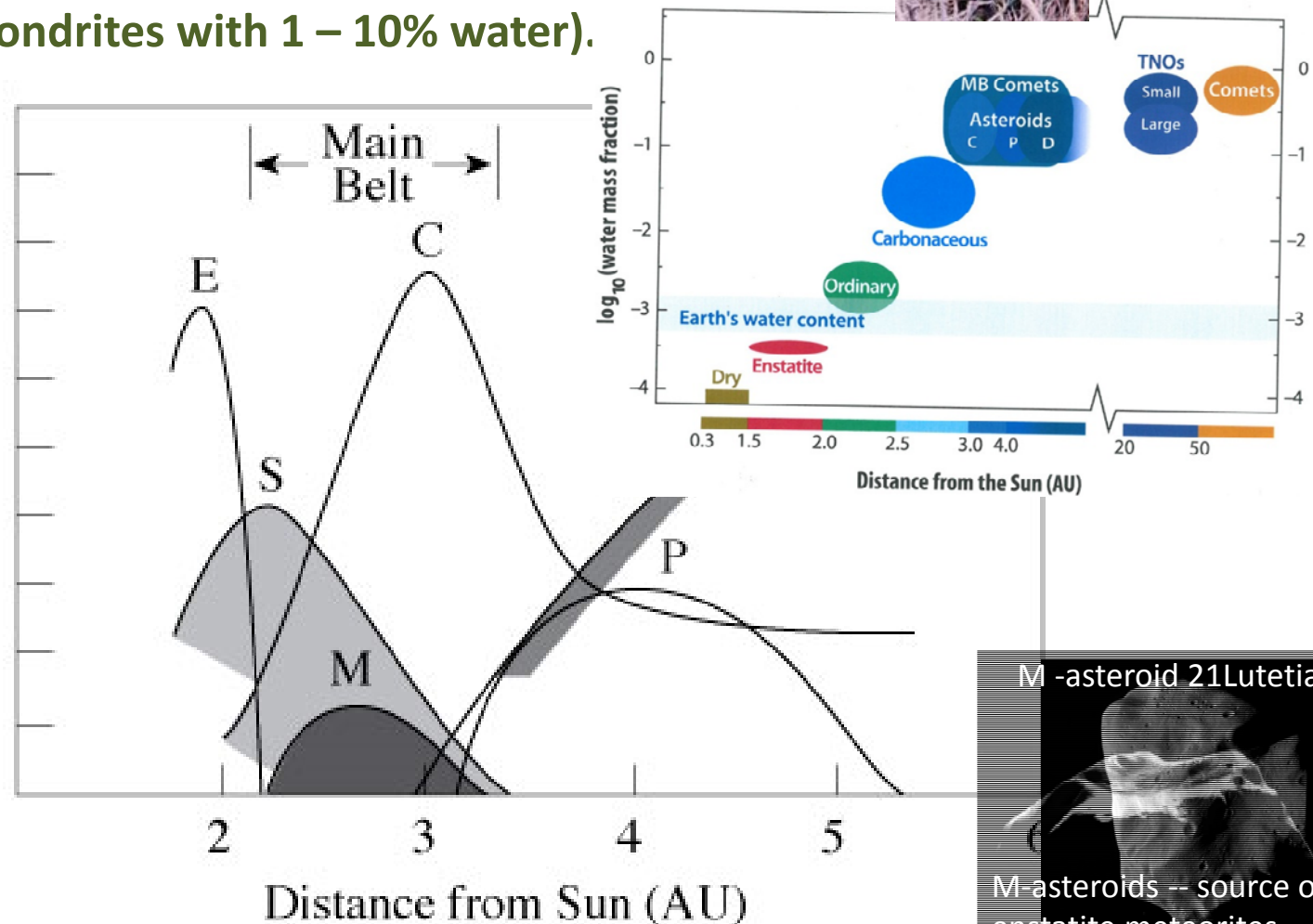
# Asteroids and water

Asteroids from the inner belt are dehydrated (e.g. S- and M-type, E-belt); Asteroids outside  $\sim 2.7$  AU contain water (e.g. C-type asteroids, represented by carbonaceous chondrites with 1 – 10% water).

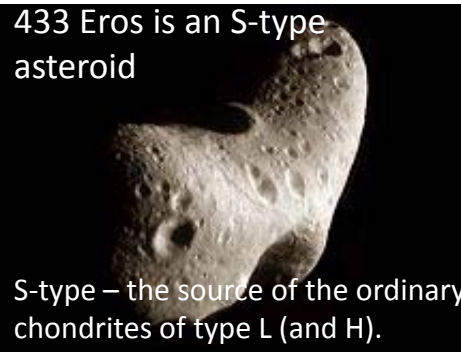
C-type asteroid  
253 Mathilde



and CM2 Maribo

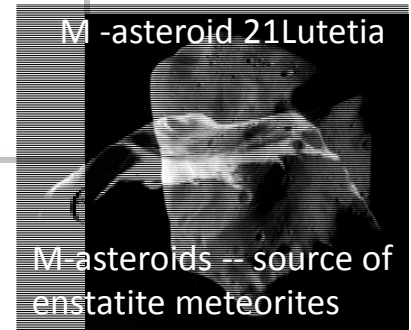


433 Eros is an S-type asteroid



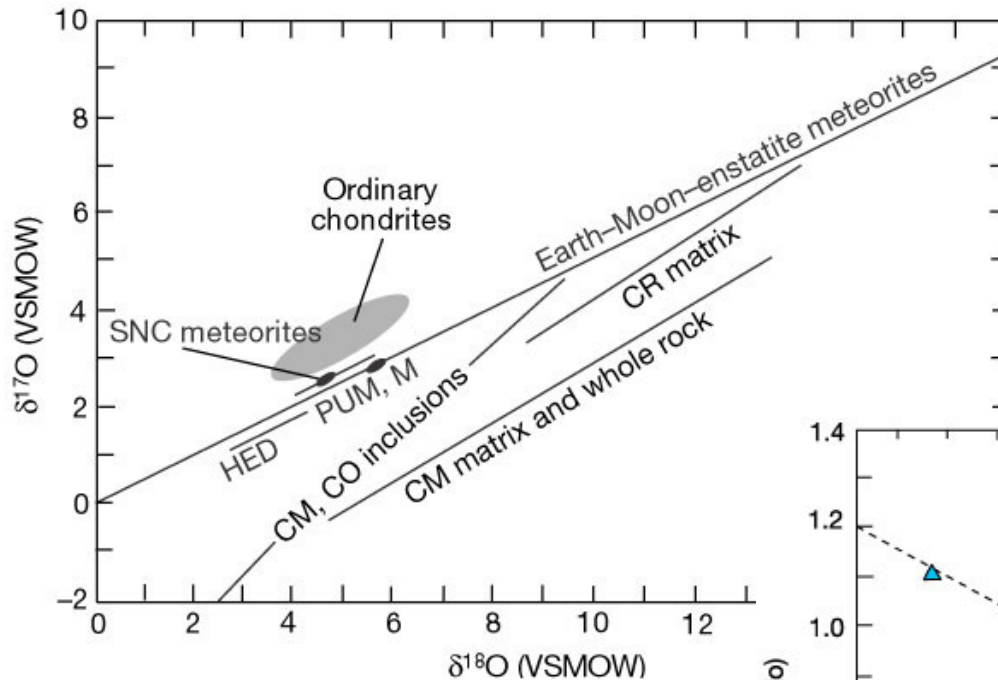
S-type – the source of the ordinary chondrites of type L (and H).

M-asteroid 21Lutetia



M-asteroids -- source of enstatite meteorites

# Bulk Earth is *not* a mixture of material condensed at many places in the protosolar nebula, but assembled locally

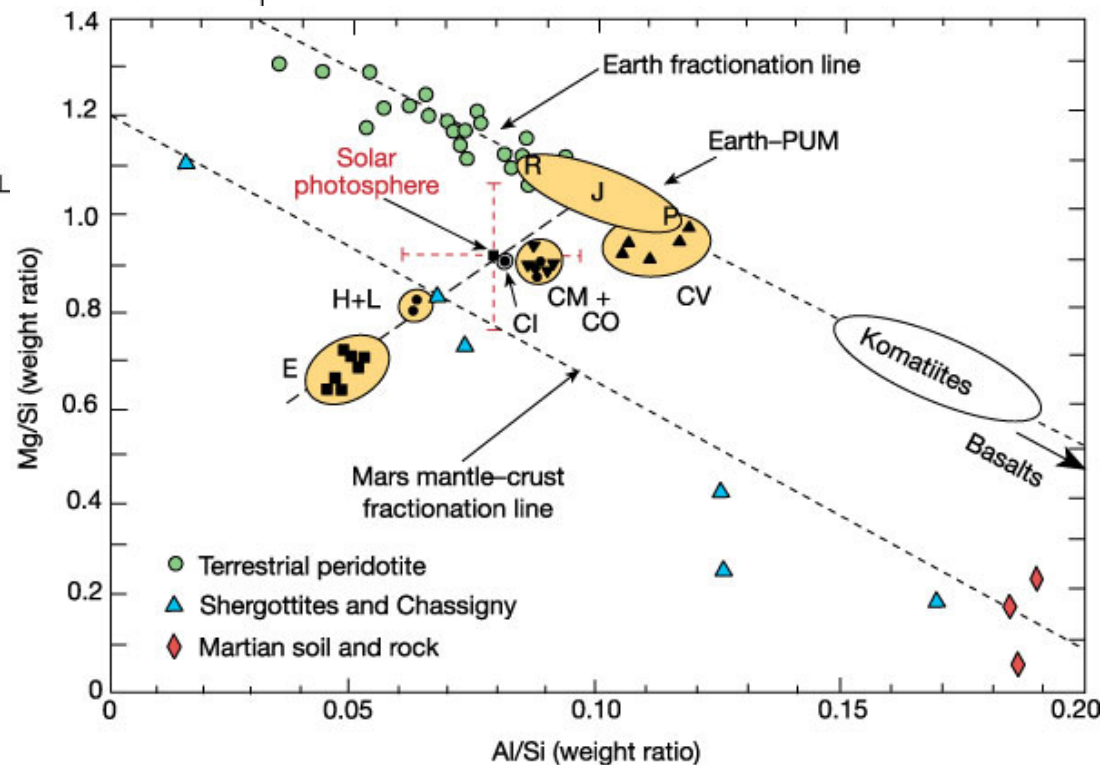


Ordinary L+H meteorites ~  
**S-type** (inner belt) asteroids  
 Enstatite E-type meteorites ~  
**M-type** (mid-belt+Mercury?) asteroids  
 Carbonaceous (CI, CM, ...) meteorites ~  
**C-type** asteroids

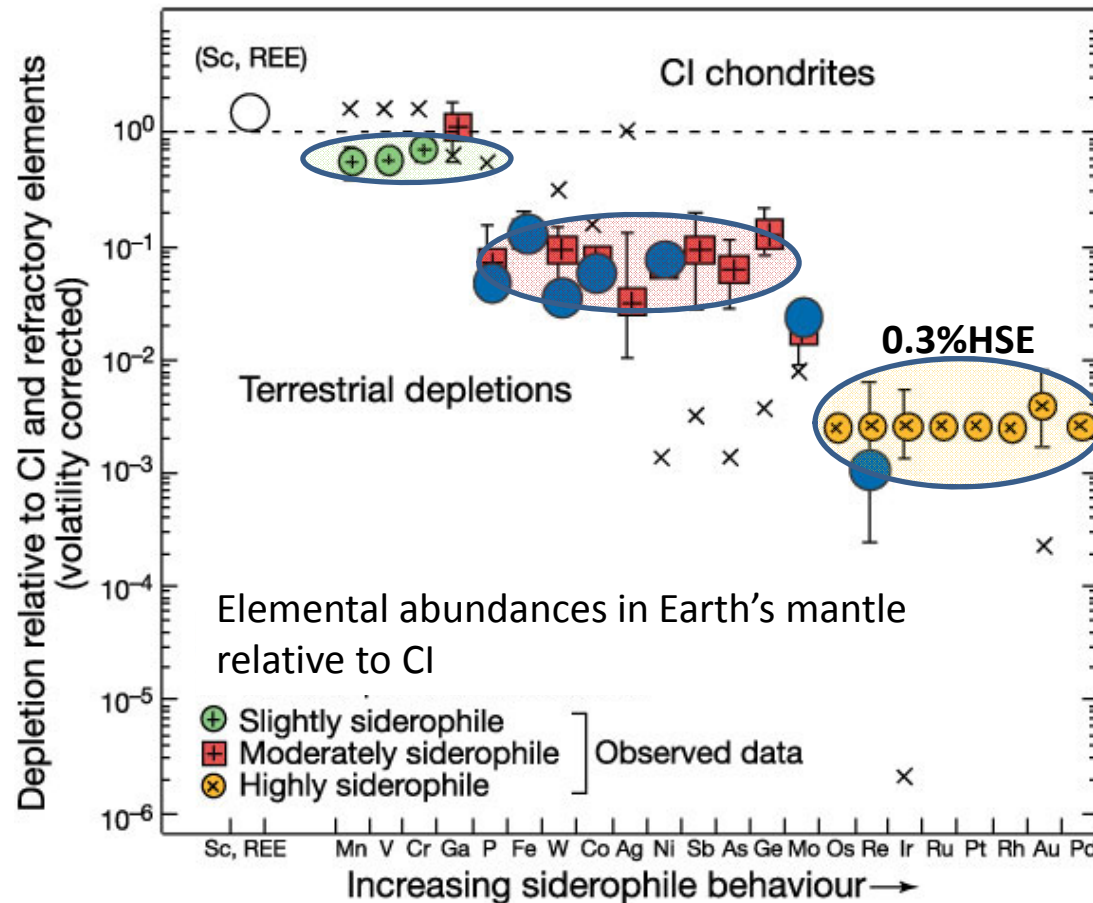
No simple mixture of meteorites  
 can explain the bulk Earth

Chemical equilibrium explain  
 gross elemental abundance  
 as function of solar distance

Oxygen isotopic ratios reflect  
 formation place;  
 They formed in Earth orbit



# The late veneer: Most likely Earth accreted 0.01 M<sub>mantle</sub> chondritic material after the core separated from the mantle



0.3% of PUM =  $3 \times 10^{21}$  kg CI  
 1 % of total mantle =  $4 \times 10^{22}$  kg  
 => 5-10% water in CI ~ from  
 $1.5 \times 10^{20}$  kg to  $4 \times 10^{21}$  kg water, or  
 from 0.1 to 3 M<sub>ocean</sub> masses  
**(uncertainty of a factor 30)**

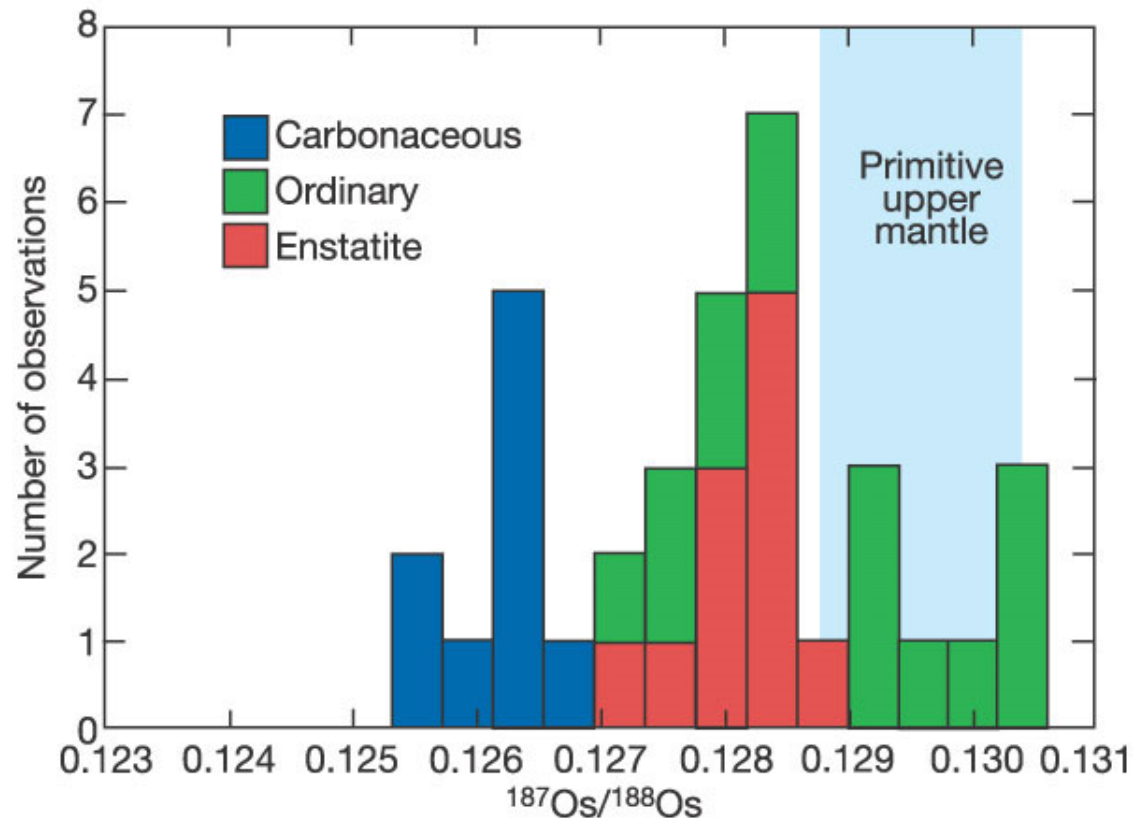
## Conclusion:

the ocean could come from  
 wet chondritic material,  
 accreted early, after the core  
 formation

Drake & Righter 2002 (Nature 416, 39),  
 see also U.Mann et al 2012

The mass of the upper mantle is  $1 \times 10^{24}$  kg, of the full silicate mantle  $4 \times 10^{24}$  kg, the crust  $2.4 \times 10^{22}$  kg,  
 the ocean  $1.4 \times 10^{21}$  kg, the full hydrosphere  $1.6 \times 10^{21}$  kg, the total Earth  $6 \times 10^{24}$  kg, the core  $1.9 \times 10^{24}$  kg  
 Primitive Upper Mantle = PUM = upper mantle plus crust ~  $1 \times 10^{24}$  kg

**$^{187}\text{Os}/^{188}\text{Os}$  ratios in carbonaceous, ordinary and enstatite chondrites, and in the Earth's primitive upper mantle (from Drake & Righter 2002).**



Core formation  
brought Os- and Ir  
down to  $\sim 5$  pg/g

If today's mantle Os + Ir  
of 3 ng/g are due to a  
late veneer, they came  
from ordinary (H-L type)  
chondrites to fit Os-isotopes

**Carbonaceous chondrites could have brought an ocean mass of water, but dry H-L type could not.**

**Cc cannot, however, explain the Os isotopes, while H-L can.**

**This may indicate that the late veneer bombardment was dry, and then ocean water came later than 0.5 Gyr into Earth's evolution.**

## If water

didn't accrete during the original assembling  
(too hot nebula, local assembly with no water, ...)

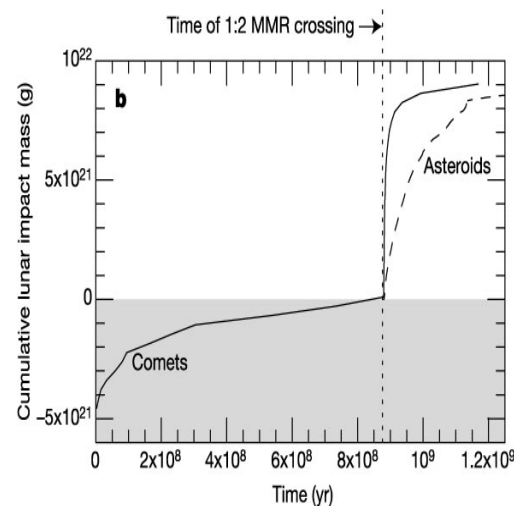
...and it  
didn't come from the late veneer (wrong Os-isotopes,...),

could it then have come during LHB ?  
-- geological evidence show existence of oceans on Earth 3.8 Gyr  
ago, and zircons may indicate some water present before that

Various runs of the Nice model code suggest different resonances and disturbances  
as the cause of the transportation of smaller bodies inward to the inner solar system  
-- some scenarios predict wet impactors, some dry impactors.

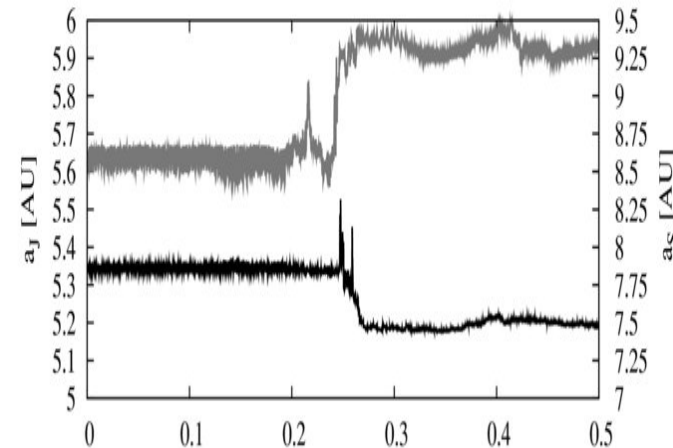
## Some of the Nice-model LHB scenaria:

Goomes et al 2005, **The 1:2 resonance scenario:**  
Jupiter and Saturn crossing 1:2 resonance during LHB, delivering 6% ocean-mass of comets and a higher mass of asteroid impacts

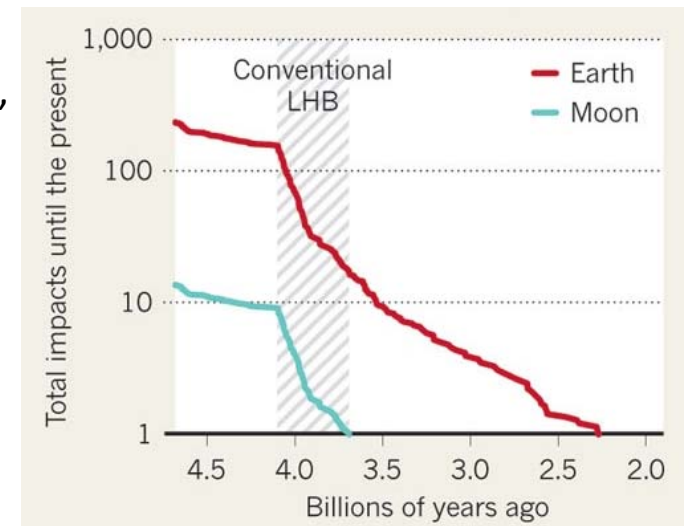


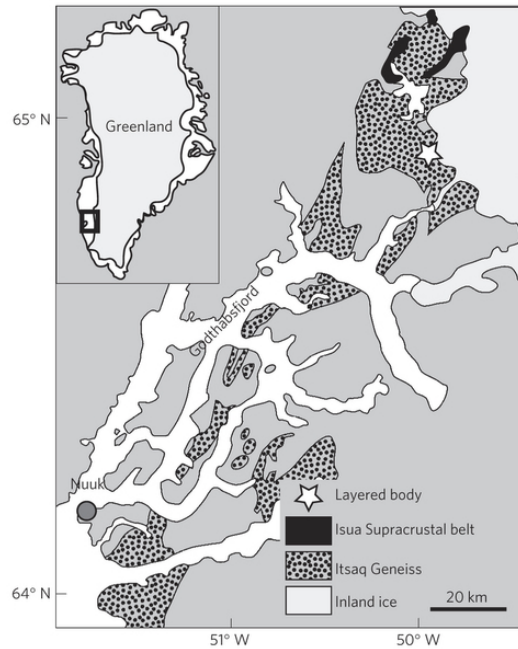
Morbidelli 2014, this meeting:  
Where are all the comets that the Nice model predicts?

Morbidelli et al 2010, **The jumping-Jupiter scenario:**  
a large cometary LHB flux ( $\sim 90\%$  of the projectiles) and few asteroid impactors, giving a wet LHB and less excentric terrestrial planet orbits than the 1:2 scenario



Bottke et al 2012,  
**The long tail bombardment :**  
a very dry LHB bombardment by E-belt asteroids starting already 4.1 Gyr ago





## Hadean crust Os indicate a substantial mantle-depletion event 4.1 Gyr ago, in agreement with the Bottke et al 2012 dynamic model of an early LHB:

Samples from the Greenland Ujaragssuit Nunat intrusion indicate a global peak in mantle depletion at  $\sim 4.1$  Gyr ago, which the authors interpret as the LHB, in agreement with an E-belt disturbance resulting in an early (and dry) LHB.

From: Coggon et al 2013, NatureGeosci 6, 871

**If water didn't come during:**

**Earth formation**

(too hot)

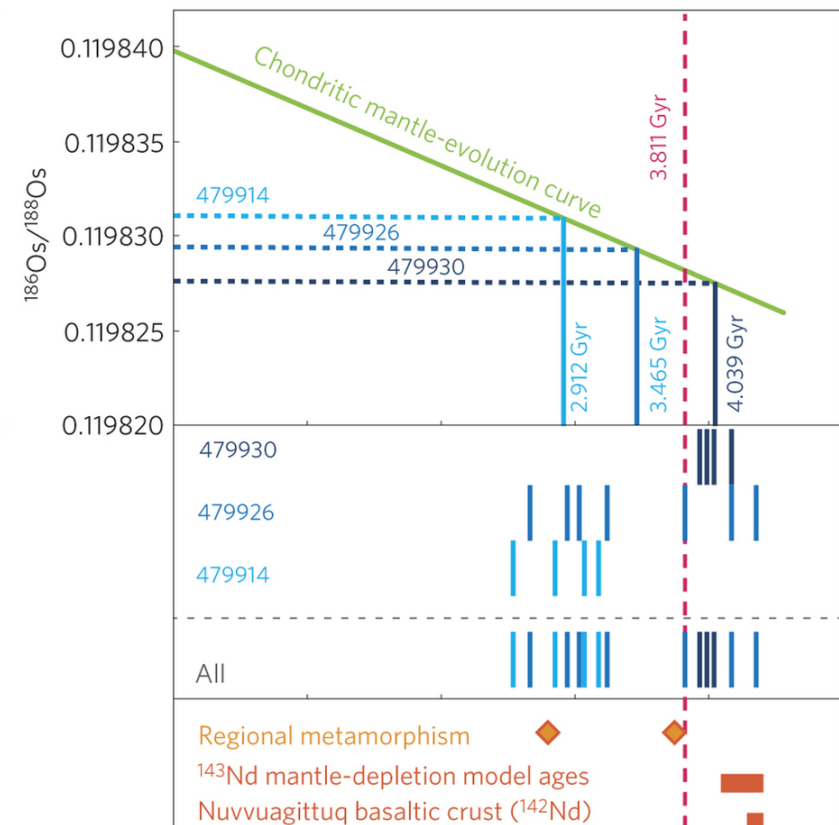
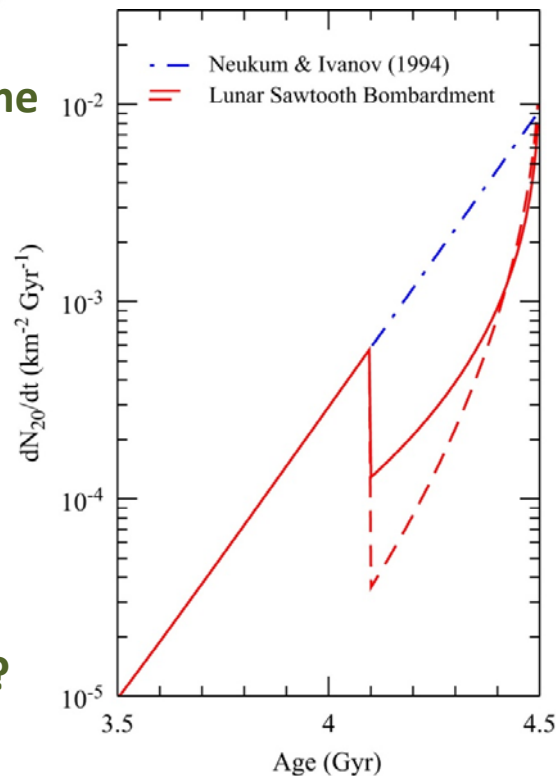
**late veneer**

(Os isotopes)

**LHB**

(E-type asteroids)

**When did it come?**



## Searches for traces of the LHB at 3.8 Gyr old Isua

In Jørgensen et al (2009) we measured 40 samples of 4 different kind of sediments with an accuracy of 2 ppt, and concluded that the cosmic contribution of Ir is small but clear (150 ppt at Isua versus 20 ppt in present-day ocean crust and upper continental crust), contradicting a chondrite-like bombardment but consistent with a cometary dominated LHB.

Koeberl et al (2000) qualitatively found the same based on fewer and more random samples and poorer resolution near the abundance level. Assuming the LHB was condrites, they concluded that they could identify no cosmic contamination at Isua.

Akilia: Anbar et al 2001 found depleted Ir (~10 ppt?) in old Akilia (150 km SW of Isua) rocks, but both age and type of the rocks were later disputed (3.85->3.65 Gyr, no sediments but rather basaltic; Moorbath 2005).

Hf-W at Isua:  $^{182}\text{Hf}(\text{litrofile}) \rightarrow ^{182}\text{W}(\text{sidrofile}) \Rightarrow$  crustal  $^{182}\text{W}/^{183}\text{W}$  on Earth is up. Schoenberg et al 2002 find 3 sigma lower  $^{182}\text{W}/^{183}\text{W}$  at Isua than present crust, which they interpret as cosmic mixing at Isua, but it should be 1000 lunar impact.

Cr at Isua:  $^{53}\text{Mn} \rightarrow ^{53}\text{Cr}$  affects the  $^{53}\text{Cr}/^{52}\text{Cr}$  ratio, but again the meteoritic value is only slightly different from the Earth's crustal value. Frey & Rossing (2005) found the Isua  $^{53}\text{Cr}/^{52}\text{Cr}$  value almost present day terrestrial, as opposed to K/T which is almost as in Allende because of the large meteoritic to terrestrial abundance ratio in K/T layers.

Moon and Earth iridium (HSE) abundance is sensitive in a different manner to cometary and asteroid bombardment

Scaling of lunar crater counts => 2000 tons per  $\text{m}^2$  on Isua

### Comets:

2000 tons comets with 50% mixing efficiency, 10% CI of 465,000 ppt Ir, mixed into 50 km crust => expected **130 ppt** in **Isua** sediments,

Moon: very low mixing => **low iridium abundance on the Moon**

### Asteroids:

2000 tons CI with 90% mixing efficiency and 465,000 ppt Ir

=> expected **3000 ppt** at 3.8 Gyr old **Isua**, Greenland,

and 50% mixing on Moon => **5000 ppt on the Moon**

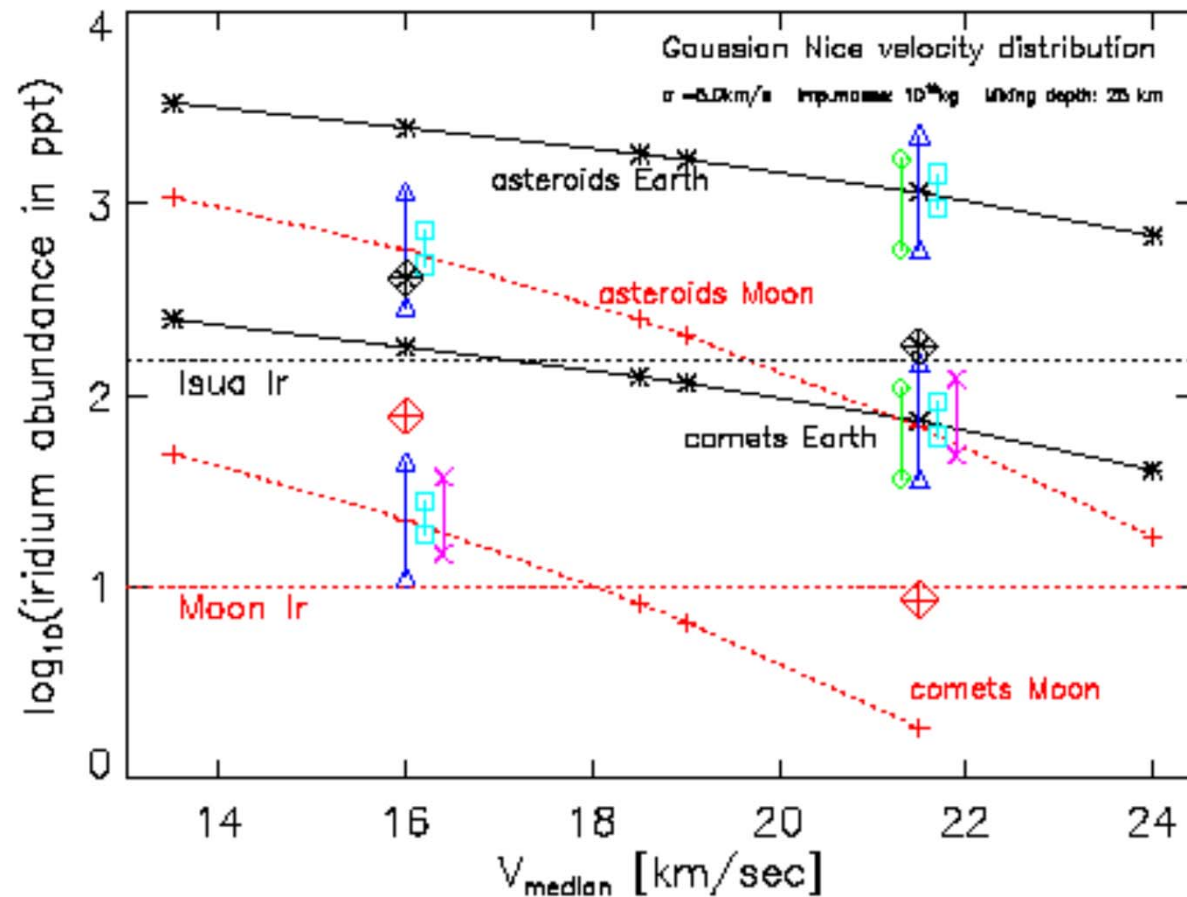


### Measurements:

**150 ppt iridium at Isua**  
**maybe low Ir in lunar rocks**  
**~expected from comets**



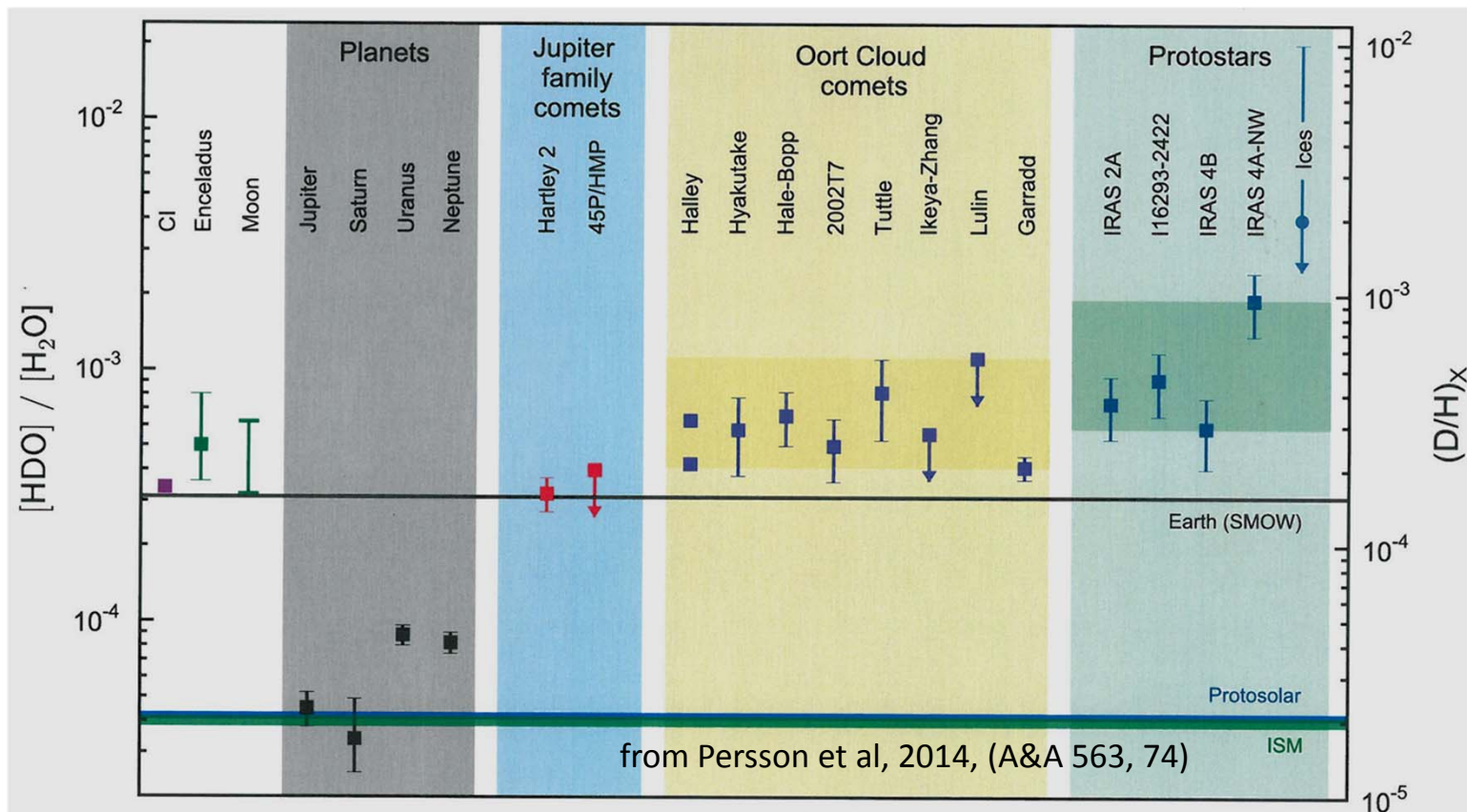
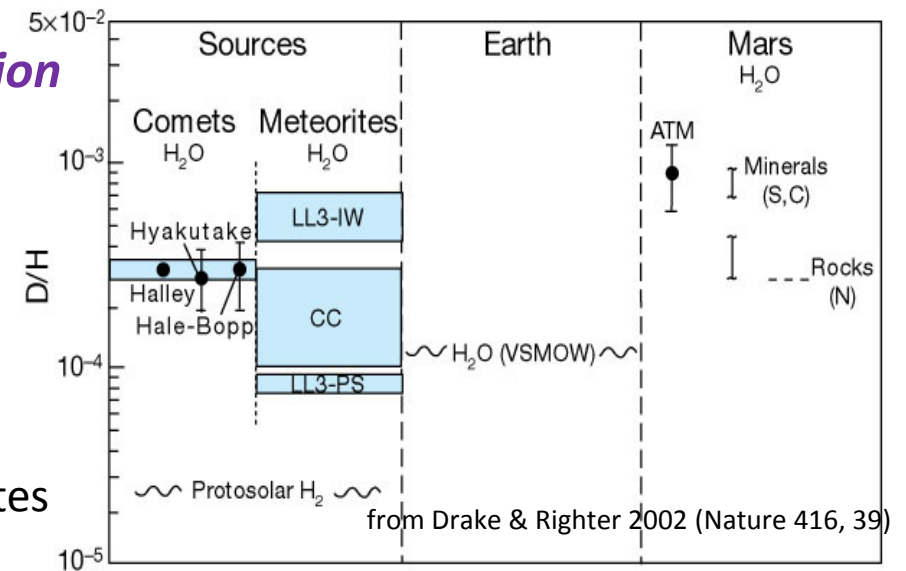
Uncertainties in inputdata are large (e.g. length of LHB, velocity of impactors, mixing depth in crust, iridium abundance in bulk lunar crust, angle dependence of plume mass), but the HSE abundance in comets, asteroids (and various meteorite types), and planetary mantles differ with such a huge factor that HSE (e.g. iridium and osmium) is one of the most sensitive indicators of impact parameters.



Formulas and assumptions given in Jørgensen et al 2009, Icarus 204, 368.

The D/H ratio in water contain *no information* about the origin of Earth's water – it could be explained by carbonaceous chondrite impactors as well as by comets.

Water in rocks from Mars has D/H as Oort cloud comets, while Earth's oceans have D/H as Jupiter family comets and carbonaceous chondrites



Sub-mm observations indicate cometary infall through young stellar debris disks, such as the beta-pictoris system. With full capability of ALMA such observations can be spatially extended, and already show D/H comparable to comets in our solar system.

